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Author(s): Justin D. R. Steinfadt, Lars Bildsten, David L. Kaplan, Benjamin J. Fulton, Steve B. Howell, T. R. Marsh, Eran O. Ofek and Avi Shporer

Article Title: A Search for Pulsations in Helium White Dwarfs

Year of publication: 2012

Link to published article:

<http://dx.doi.org/10.1086/663865>

Publisher statement: © 2012. The Astronomical Society of the Pacific.



CHICAGO JOURNALS



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Reviewed work(s):

Source: *Publications of the Astronomical Society of the Pacific*, Vol. 124, No. 911 (January 2012), pp. 1-13

Published by: [The University of Chicago Press](#) on behalf of the [Astronomical Society of the Pacific](#)

Stable URL: <http://www.jstor.org/stable/10.1086/663865>

Accessed: 13/03/2012 07:08

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A Search for Pulsations in Helium White Dwarfs

JUSTIN D. R. STEINFADT,¹ LARS BILDSTEN,^{1,2} DAVID L. KAPLAN,³ BENJAMIN J. FULTON,⁴
 STEVE B. HOWELL,⁵ T. R. MARSH,⁶ ERAN O. OFEK,^{7,8} AND AVI SHPORER^{1,4}

Received 2011 May 1; accepted 2011 November 17; published 2012 December 19

ABSTRACT. The recent plethora of sky surveys, especially the Sloan Digital Sky Survey, have discovered many low-mass ($M < 0.45 M_{\odot}$) white dwarfs that should have cores made of nearly pure helium. These WDs come in two varieties: those with masses $0.2 < M < 0.45 M_{\odot}$ and H envelopes so thin that they rapidly cool and those with $M < 0.2 M_{\odot}$ (often called extremely low mass [ELM] WDs) that have thick enough H envelopes to sustain 10^9 yr of H burning. In both cases, these WDs evolve through the ZZ Ceti instability strip, $T_{\text{eff}} \approx 9000\text{--}12,000$ K, where g -mode pulsations always occur in carbon/oxygen WDs. This expectation, plus theoretical work on the contrasts between C/O and He-core WDs, motivated our search for pulsations in 12 well-characterized helium WDs. We report here on our failure to find any pulsators among our sample. Though we have varying amplitude limits, it appears likely that the theoretical expectations regarding the onset of pulsations in these objects require closer consideration. We close by encouraging additional observations as new He WD samples become available, and we speculate on where theoretical work may be needed.

1. INTRODUCTION

White dwarf asteroseismology offers the opportunity of directly constraining mass, core composition, envelope composition and stratification, spin rate, and magnetic field strength (Castanheira & Kepler 2008). The ZZ Ceti H-atmosphere white dwarfs (WDs) pulsate in normal modes of oscillation (g modes) where buoyancy is the restoring force (Winget & Kepler 2008). These stars exhibit pulsations when they enter the instability region, a discrete strip in the $T_{\text{eff}}\text{--}\log g$ plane that spans $11,000 \lesssim T_{\text{eff}} \lesssim 12,250$ K at $\log g = 8.0$. This strip has been extensively observed and constrained by numerous sources (Wesemael et al. 1991; Mukadam et al. 2004; Castanheira et al. 2007; Gianninas et al. 2007) and is likely pure; i.e., all cooling WDs pulsate as they evolve through the instability strip. These data offer insight into the evolution of WDs (Fontaine et al. 1982; Cassisi et al. 2007). The lowest-mass WD known to pulsate is HE 0031-5525 at $\log g = 7.65$ and $T_{\text{eff}} = 11,480$ K and

a mass of $\approx 0.5 M_{\odot}$ (Castanheira et al. 2006). The known ZZ Ceti pulsators are certainly C/O-core WDs, where the theoretical understanding has been focused (Winget et al. 1982b; Brassard et al. 1991; Brassard & Fontaine 1997; Winget & Kepler 2008).

Starting with Bergeron et al. (1992) and Marsh et al. (1995), there has been a growing realization that many WDs have $\log g \lesssim 7.76$, implying core masses lower than expected for C/O-core WDs. These are expected to be He WDs made from the truncated red giant branch (RGB) evolution of $M < 2.3 M_{\odot}$ stars (D’Cruz et al. 1996; Dominguez et al. 1999; Pietrinferni et al. 2004; VandenBerg et al. 2006; Panei et al. 2007). To form a He WD, a process removes mass from the H envelope, quenching the H-shell burning that grows the He core to ignition. Two modes of evolution have been proposed to truncate RGB evolution: mass loss due to binary interaction and mass loss due to stellar winds. As a progenitor star ascends the RGB, its radius increases rapidly, eventually filling the Roche lobe, and mass transfer begins. If unstable, this mass transfer initiates a common envelope event (Paczynski 1976; Iben & Livio 1993), ejecting the envelope from the red giant as the He core and companion spiral into shorter orbital periods (Taam & Sandquist 2000; Deloye & Taam 2010; Ge et al. 2010). It is this physics that leads to the expectation that He WDs should be in binary systems (Marsh et al. 1995; Kilic et al. 2007b; Agüeros et al. 2009). However, surveys have uncovered significant numbers of He WDs that do not appear to be in binary systems (Agüeros et al. 2009; Brown et al. 2010, 2011a). In stars of high metallicity, the increased opacity in the envelope due to metals can drive very strong stellar winds that cause mass loss in excess of $0.5 M_{\odot}$ during the RGB phase (Han et al. 1994; Lee et al. 1994;

¹ Department of Physics, Broida Hall, University of California, Santa Barbara, CA 93106; jdrsteinfadt@gmail.com.

² Kavli Institute for Theoretical Physics and Department of Physics, Kohn Hall, University of California, Santa Barbara, CA 93106; bildsten@kitp.ucsb.edu.

³ Physics Department, University of Wisconsin - Milwaukee, Milwaukee WI 53211; kaplan@uwm.edu.

⁴ Las Cumbres Observatory Global Telescope, 6740 Cortona Drive, Suite 102, Goleta, CA 93117; bjfulton@lco.net, ashporer@lco.net.

⁵ NASA Ames Research Center.

⁶ Department of Physics, University of Warwick, Coventry CV4 7AL, UK.

⁷ Division of Physics, Mathematics and Astronomy, California Institute of Technology, Pasadena, CA 91125.

⁸ Einstein Fellow.

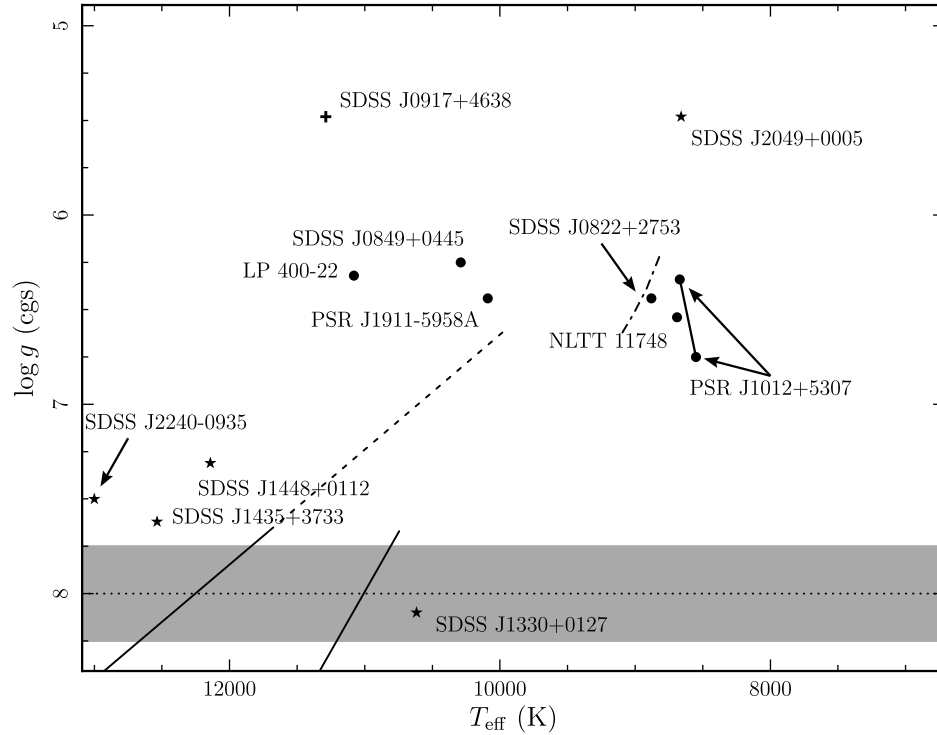


FIG. 1.—The $\log g$ - T_{eff} plane. The dark-gray horizontal strip illustrates the location of the C/O WDs. The solid lines represent the fitted boundaries of the empirical ZZ Ceti instability strip as determined by Gianninas et al. (2007) with the dashed line merely being the extension of the blue edge. The dash-dotted line represents the theoretical blue edge highlighted for very low mass He WDs by Steinfadt et al. (2010a). We plot only objects we have observed. SDSS J1330 + 0127, SDSS J1435 + 3733, SDSS J1448 + 0112, SDSS J2049 + 0005, and SDSS J2240-0935, the stars, are from Eisenstein et al. (2006) (corrected by Silvestri et al. 2006 and Heller et al. 2009). SDSS J0822 + 2753 and SDSS J0849 + 0445 are from Kilic et al. (2010b). SDSS J0917 + 4638, the plus, is from Kilic et al. (2007a). NLTT 11748 is from Kawka & Vennes (2009), Steinfadt et al. (2010b), and Kilic et al. (2010a). LP 400-22 is from Kawka et al. (2006). PSR J1012 + 5307 is from two sources, neither obviously superior, van Kerkwijk et al. (1996) and Callanan et al. (1998). PSR J1911-5958A is from Bassa et al. (2006a).

TABLE 1
OBSERVED CANDIDATES AND RESULTS

Object	SDSS g' (mag)	$\log g$ (cgs)	T_{eff} (K)	Reference	Det. limit ^a (mmag)
SDSS J0822+2743	18.3	6.55 ± 0.11	8880 ± 60	Kilic et al. (2010b)	8
SDSS J0849+0445	19.3	6.23 ± 0.08	$10,290 \pm 250$	Kilic et al. (2010b)	15
SDSS J0917+4638	18.7	5.48 ± 0.03	$11,288 \pm 72$	Kilic et al. (2007a)	11
SDSS J1330+0127	18.9	8.1 ± 0.2	$10,617 \pm 452$	Silvestri et al. (2006)	44
SDSS J1435+3733	17.1	7.62 ± 0.12	$12,536 \pm 488$	Eisenstein et al. (2006)	4
SDSS J1448+0112	19.6	7.31 ± 0.18	$12,142 \pm 486$	Eisenstein et al. (2006)	34
SDSS J2049+0005	19.7	5.48 ± 0.10	8660 ± 144	Eisenstein et al. (2006) ^b	21
SDSS J2240-0935	17.6	6.96 ± 0.09	$11,449 \pm 205$	Heller et al. (2009)	14
LP 400-22	17.2	6.32 ± 0.08	$11,080 \pm 140$	Kawka et al. (2006)	4
NLTT 11748	17.1 ^c	6.54 ± 0.05	8690 ± 140	Kilic et al. (2010a)	5
PSR J1012+5307	19.5 ^d	6.34 ± 0.20	8670 ± 300	Callanan et al. (1998)	20
		6.75 ± 0.07	8550 ± 25	van Kerkwijk et al. (1996)	...
PSR J1911-5958A	22.2 ^d	6.44 ± 0.20	$10,090 \pm 150$	Bassa et al. (2006a)	16

^aReporting the best limit over a minimum of 1–7 mHz frequency range (where possible). See Table 3 for frequency range details and multiple epochs.

^bSDSS J2049 + 0005 is likely a distant A star (Eisenstein et al. 2006; Kilic et al. 2007a).

^cMagnitude in Johnson B .

^dMagnitude in Johnson V .

TABLE 2
DETAILS OF OBSERVATIONS

Object	SDSS g' (mag)	Telescope/camera/filter	Date (yyyymmdd)	Length (hr)	Exp. time (s)	No. of exp. ^a
SDSS J0822+2743	18.3	WHT/ACAM/ g'	20110220	2.1	30	137
SDSS J0849+0445	19.3	WHT/ACAM/ g'	20110117	1.9	60	96
SDSS J0917+4638	18.7	WHT/ACAM/ g'	20110220	2.3	45	145
SDSS J1330+0127	18.9	WIYN/OPT/BG39	20070530	1.6	90	57
SDSS J1435+3733	17.1	WIYN/OPT/BG39	20070529	1.5	15	197 ^b
		WIYN/OPT/BG39	20070531	1.9	15	233 ^b
		WIYN/OPT/BG39	20070601	3.8	45	259 ^b
SDSS J1448+0112	19.6	WIYN/OPT/BG39	20070531	2.0	60	104
SDSS J2049+0005	19.7	WIYN/OPT/BG39	20070531	3.4	60	166
SDSS J2240-0935	17.6	P60/P60CCD/ g	20060804	3.7	30	204
		P60/P60CCD/ g	20060817	4.5	100	73 ^c
		P60/P60CCD/ g	20060919	1.1	30	57
		P60/P60CCD/ g	20060925	1.0	30	60
		P60/P60CCD/ g	20061205	0.9	30	59
LP 400-22	17.2	WIYN/OPTIC/ g'	20070530	2.4	20	269
NLTT 11748	17.1 ^d	FTN/Merope/ g'	20091222	4.0	45	198 ^b
		FTN/Merope/ g'	20100108	3.3	45	167 ^b
PSR J1012+5307	19.5 ^e	WHT/ACAM/ g'	20100330	2.0	60	103
PSR J1911-5958A	22.2 ^d	HST/WFC3/F200LP	20100305	6.0	60	142

^a Excludes images contaminated by cosmic rays, clouds, and excess noise.

^b Also excludes images taken during eclipse.

^c Observation contains large gaps.

^d Magnitude in Johnson B .

^e Magnitude in Johnson V .

TABLE 3
RESULTS OF OBSERVATIONS

Object	Date	Filter	Detection limit (mmag)	Frequency range (mHz)	Figure
SDSS J0822+2743	2011/02/20	SDSS g'	8	0.5–11	2
SDSS J0849+0445	2011/01/17	SDSS g'	15	0.3–7	3
SDSS J0917+4638	2011/02/20	SDSS g'	11	0.6–8	4
SDSS J1330+0127	2007/05/30	BG-39	44	0.3–5	5
SDSS J1435+3733	2007/05/29	BG-39	4	0.9–17	6
	2007/05/31	BG-39	6	0.9–17	6
	2007/06/01	BG-39	4	0.3–9	6
SDSS J1448+0112	2007/05/31	BG-39	34	0.5–7	7
SDSS J2049+0005	2007/05/31	BG-39	21	0.4–7	8
SDSS J2240-0935	2006/08/04	Gunn g	22	0.3–7	9
	2006/08/17	Gunn g	13	0.3–2.25	9
	2006/09/19	Gunn g	14	0.6–7	9
	2006/09/25	Gunn g	16	0.7–9	9
	2006/12/05	Gunn g	48	0.5–9	10
LP 400-22	2007/05/30	SDSS g'	4	0.5–15	11
NLTT 11748	2009/12/22	SDSS g'	5	0.2–7	12
	2010/01/08	SDSS g'	4	0.2–7	12
PSR J1012+5307	2010/03/30	SDSS g'	20	0.2–7	13
PSR J1911-5958A	2010/03/05	F200LP	16	0.05–3.25	13

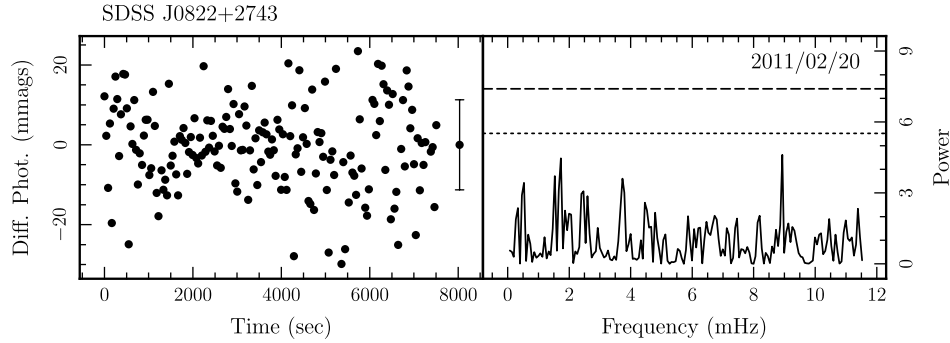


FIG. 2.—Light curve and Lomb-Scargle periodogram for SDSS J0822 + 2743 for observations taken at WHT/ACAM on 2011 February 20. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

Catelan 2000; Hansen 2005). Therefore, it is possible that high-metallicity stars may lose enough mass on the RGB due to stellar winds to avoid He-core ignition and thus leave $M \gtrsim 0.4 M_{\odot}$ He WDs as remnants. Direct evidence for this scenario is scarce, but may exist in the overluminous low-mass He WDs in the globular cluster NGC 6791 (Hansen 2005; Kalirai et al. 2007; García-Berro et al. 2010). It is clear that this wind mass-loss scenario cannot produce low-mass He WDs.

Once formed, He WD evolution is different from that of C/O WDs, because nuclear burning may still play a significant role, especially when $M \lesssim 0.2 M_{\odot}$. Studies of millisecond pulsars suggest that many He WDs are born with thick ($\sim 10^{-3}$ – $10^{-2} M_{\odot}$) H envelopes (Alberts et al. 1996; van Kerkwijk et al. 2005; Bassa et al. 2006a, 2006b). These thick H envelopes develop long diffusive tails into the He core, where the higher densities and temperatures can cause them to ignite (Driebe et al. 1999; Serenelli et al. 2002; Panei et al. 2007). This ignition highlights a dichotomy in the evolution of He WDs. For He WDs of $\gtrsim 0.18$ – $0.20 M_{\odot}$ (dependent upon metallicity) the H envelopes are radially thin enough that thin-shell instabilities develop and a series of rapid thermonuclear flashes occur. These flashes burn away the envelope, leaving only a thin layer of H, which allows for rapid cooling on timescales of tens to hundreds

of millions of years. For He WDs of $\lesssim 0.18$ – $0.20 M_{\odot}$, the nuclear burning is stable and provides a dominant energy source that slows the cooling evolution of the WD to timescales of several billions of years (Serenelli et al. 2002; Panei et al. 2007). From an observational standpoint, very low mass He WDs ($< 0.2 M_{\odot}$) may spend significantly more time in a ZZ Ceti-like instability strip than the higher-mass He WDs. However, no systematic search for their pulsations has been undertaken. This is the goal of our work.

Some differences from the C/O-core ZZ Ceti stars are expected. Steinfadt et al. (2010b) used very low mass He WD models ($< 0.2 M_{\odot}$) with stable H-burning envelopes and the Brickhill convective driving criterion (Brickhill 1983, 1991; Wu & Goldreich 1999) to locate He WD pulsation candidates. The result of this work is seen in the “blue” edge boundary depicted between $\log g = 6.0$ and 6.5 in Figure 1. Clearly, many theoretical issues remain in pinpointing the precise location of the blue edge, such as the effects of convective efficiency and the use of more detailed fully nonadiabatic calculations, which may move this boundary by several hundred kelvins (Winget et al. 1982a; Fontaine & Brassard 2008). For normal-mass He WDs (0.2 – $0.45 M_{\odot}$), the blue edge of the instability strip has not been theoretically determined (see

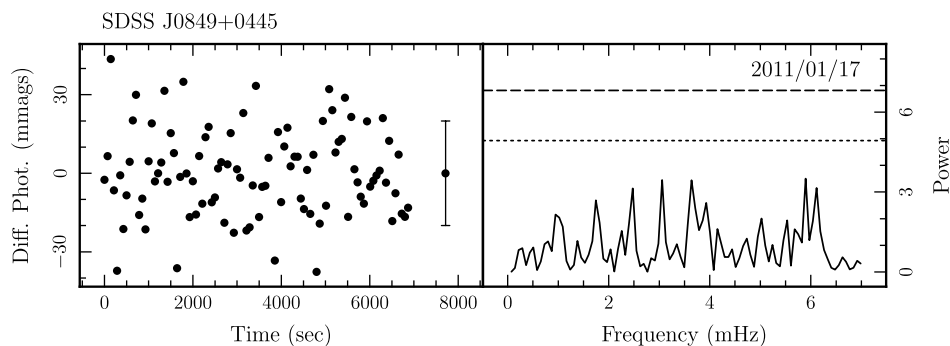


FIG. 3.—Light curve and Lomb-Scargle periodogram for SDSS J0849 + 0445 for observations taken at WHT/ACAM on 2011 January 17. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

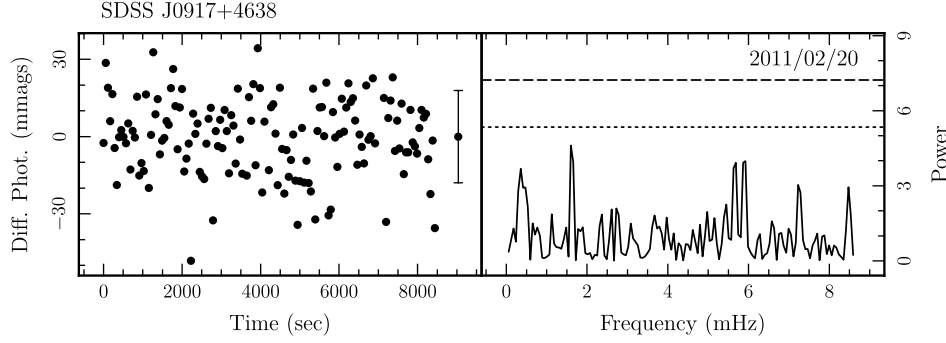


FIG. 4.—Light curve and Lomb-Scargle periodogram for SDSS J0917 + 4638 for observations taken at WHT/ACAM on 2011 February 20. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

Arras et al. 2006 for some preliminary calculations), and so we naively extrapolate the empirical boundary of the ZZ Ceti stars as determined by Gianninas et al. (2007) and connect it to the very low mass He WD boundary. Given the many theoretical uncertainties and naive extrapolations, the boundary depicted in Figure 1 is meant only to highlight the parameter space in which to carry out our search.

In this article we present our search for He WD pulsators. In § 2, we describe our selection criteria for observing candidate pulsators using objects culled from a multitude of sources. In § 3, we describe our observational strategy and analysis procedures. In § 4, we present our observations of 12 candidate objects and our detection limits. No significant pulsations have been detected; however, our detection limits may provide interesting constraints on theory, as we describe in § 5.

2. DETERMINING THE HE WD TARGETS

To select He WD candidates for observation, we generally select only those WDs with $\log g < 7.6$ (a mass $< 0.45 M_{\odot}$) and $7000 < T_{\text{eff}} < 12,000$ K. We select a larger range of temperatures than suggested by ZZ Ceti properties (and the theoretical analysis of Steinfadt et al. 2010a), as nondetections will constrain the parameter space in which the conventional theory

must hold. Of added interest is the potential detection of pulsations in parameter space that are not predicted by theory.

We find most of our candidates from surveys looking for other phenomena, e.g., the Sloan Digital Sky Survey (SDSS), but we have also found candidates in single-object articles. Companions to pulsars have been found to be good candidates, such as PSR J1911-5958A (Bassa et al. 2006a) and PSR J1012 + 5307 (van Kerkwijk et al. 1996; Callanan et al. 1998).

Eisenstein et al. (2006) offer a catalog of WDs within the SDSS that is the largest to date. They found their WDs by selecting SDSS spectra using a series of color cuts and then fitting to a large grid of WD atmospheric models. Objects with spectral contamination due to a companion star, e.g., an M-dwarf star, were not decontaminated and thus the spectral fits obtained were often erroneous. This is an issue for He WDs, as their primary formation channel requires a binary companion, which frequently contaminates the spectrum. Many of these contaminated spectra were reanalyzed by Silvestri et al. (2006) and Heller et al. (2009) in two catalogs of close binary systems. In these cases, the derived parameters of Silvestri et al. (2006) and Heller et al. (2009) were used. Since the SDSS sample did not target WDs specifically, the quality of the spectra of many of the WDs is low, as are the model fits derived from them.

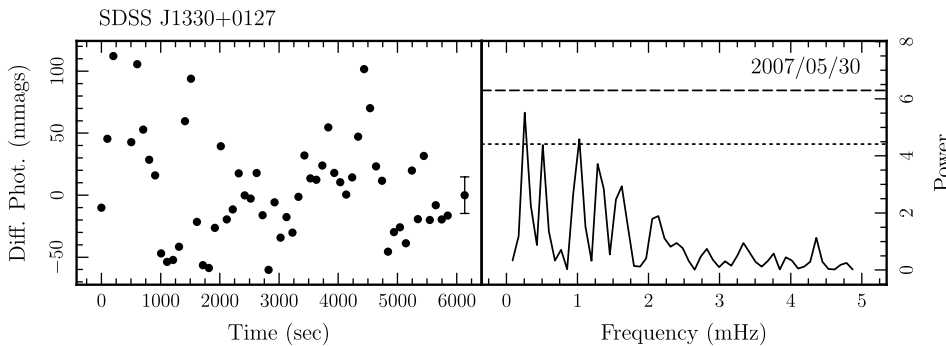


FIG. 5.—Light curve and Lomb-Scargle periodogram for SDSS J1330 + 0127 for observations taken at WIYN/OPTIC on 2007 May 30. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

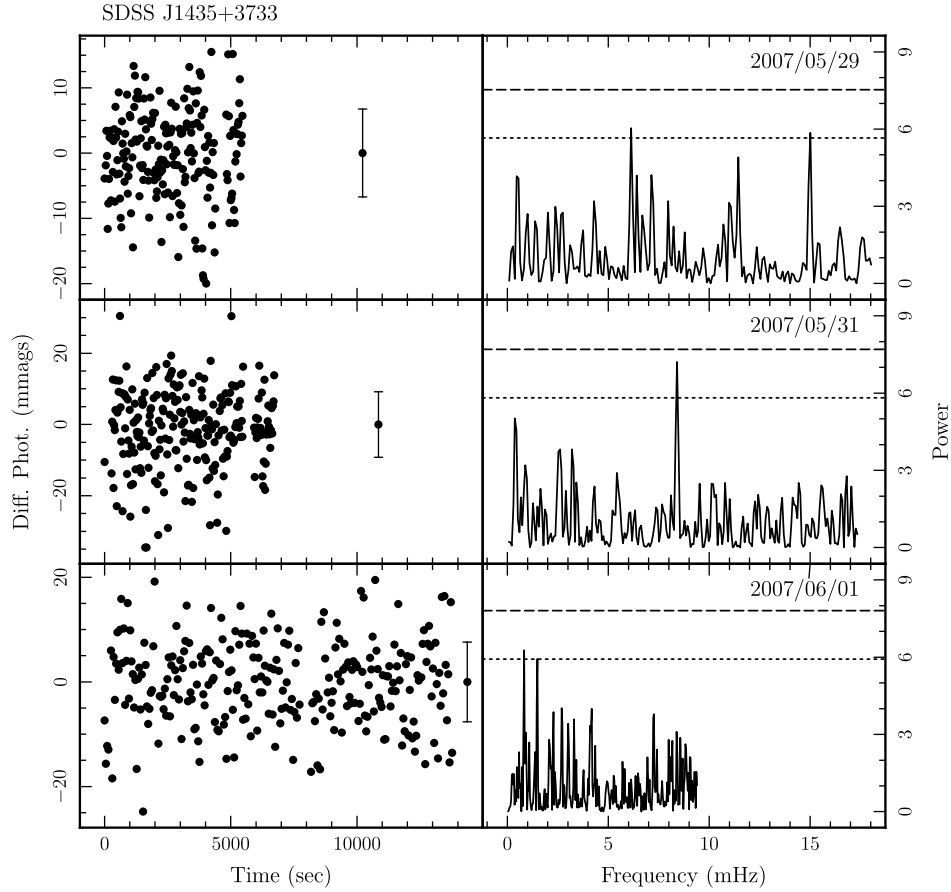


FIG. 6.—Light curves and Lomb-Scargle periodograms for SDSS J1435 + 3733 for observations taken at WIYN/OPTIC on 2007 May 29, 31, and June 1. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

Therefore, we generally do not pursue SDSS targets unless additional spectra have been taken. A few SDSS objects were observed early in our program, before this contamination was well understood.

Brown et al. (2006) performed a hypervelocity star survey looking for B-type stars escaping the Galaxy. They used SDSS

photometry to color-select their B-type stars and gathered high signal-to-noise ratio (S/N) spectra to search for large radial velocities. Given their color selection, Kilic et al. (2007a, 2007b, 2010b) made great use of this data set, uncovering many low-mass He WDs. Kawka et al. (2006) and Kawka & Vennes (2009) have used a technique similar to the Brown et al. (2006) survey

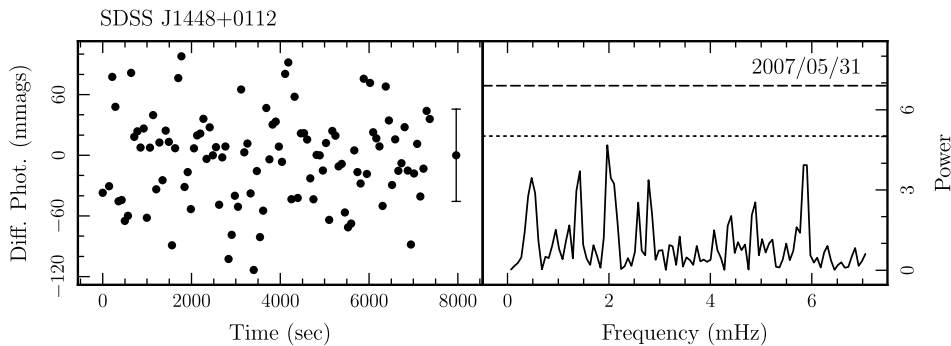


FIG. 7.—Light curve and Lomb-Scargle periodogram for SDSS J1448 + 0112 for observations taken at WIYN/OPTIC on 2007 May 31. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

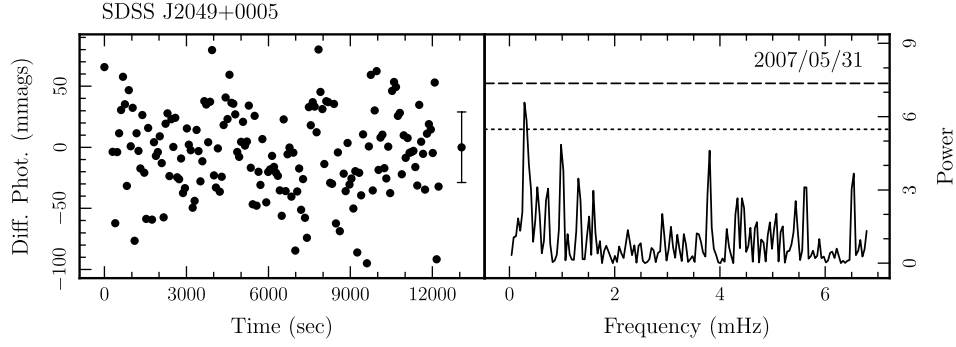


FIG. 8.—Light curve and Lomb-Scargle periodogram for SDSS J2049 + 0005 for observations taken at WIYN/OPTIC on 2007 May 31. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

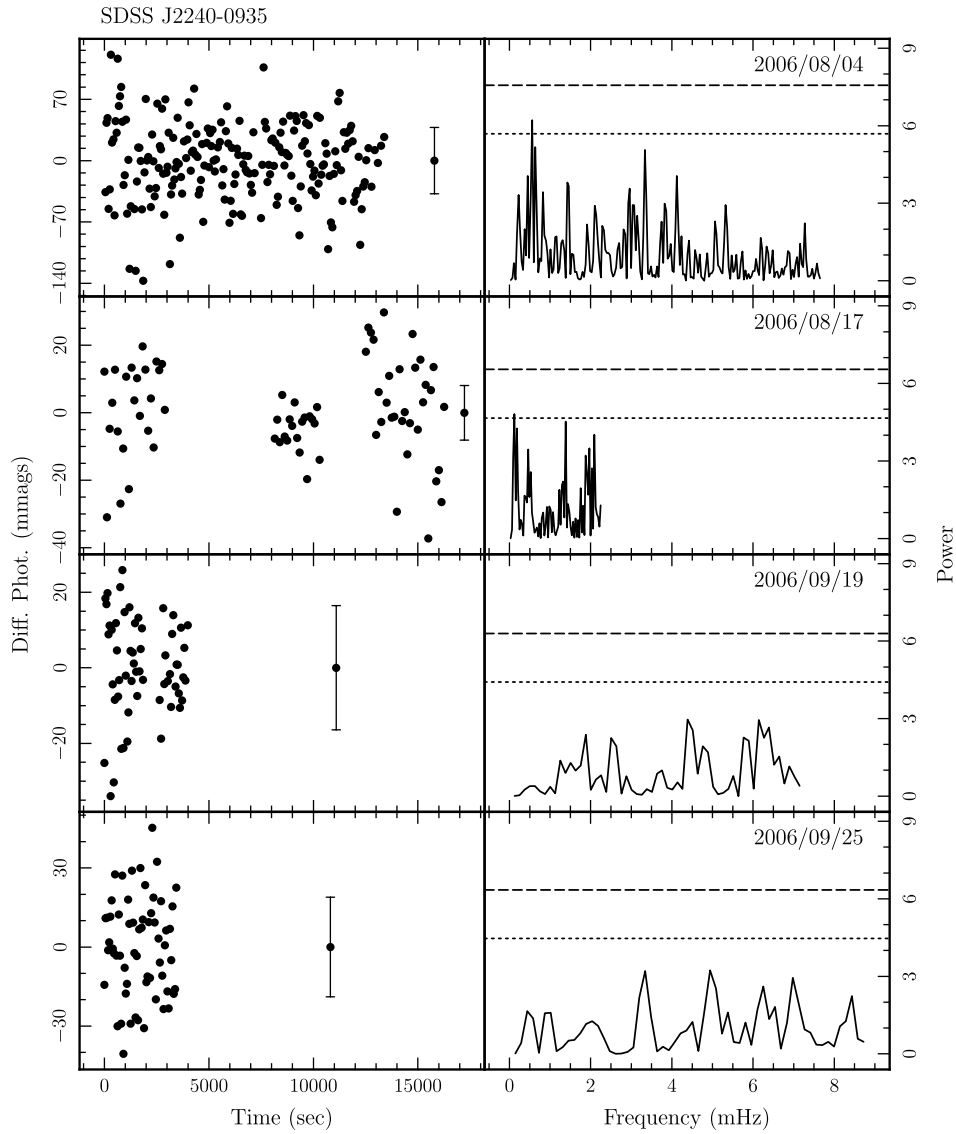


FIG. 9.—Light curves and Lomb-Scargle periodograms for SDSS J2240-0935 for observations taken at P60/P60CCD on 2006 August 4 and 17 and September 19 and 25. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

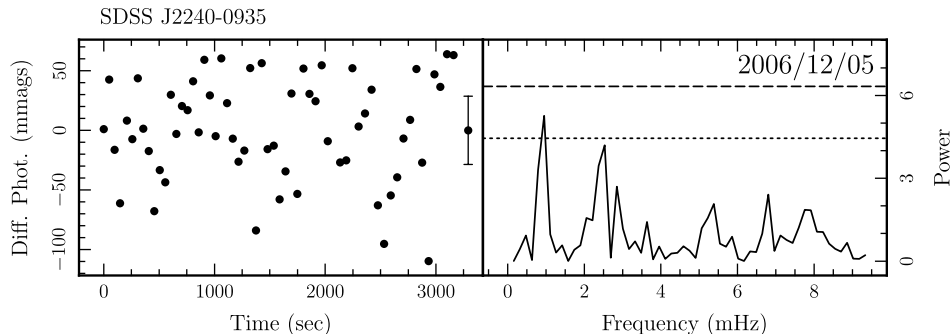


FIG. 10.—Continuation of Fig. 9 for observations taken on 2006 December 5.

and searched for hypervelocity stars in proper-motion data from a variety of sources (Oswalt et al. 1993, New Luyten Two-Tenths [NLTT], Proper Motion Survey). Their survey has discovered some of the lowest-mass WDs known (Kawka et al. 2006; Kawka & Vennes 2009). These surveys revealed dozens of candidates, only 12 of which are observed (see Table 1). Of these 12 objects, 10 are potential He WDs and seven are very low mass He WDs.

3. OBSERVATIONAL STRATEGY

The observations we executed followed a simple strategy. Steinfadt et al. (2010a) showed that the expected periods of pulsations range between 250 and 1000 s. Extrapolating our knowledge of the ZZ Ceti stars, we considered a fiducial amplitude of 10 mmag. We designed each observation to be sensitive to these properties by incorporating a fast cadence (usually less than 60 s) and individual measurements of the target that met or exceeded 10 mmag in precision (in the Poisson limit, 10,000 photons per image). In Table 2 we report on the details of our observations.

3.1. Aperture Photometry

All images were reduced using standard methods (bias-subtracted and flat-fielded). In most cases, the CCD detectors undersampled the point-spread function (PSF) of the stars, making PSF-fitting photometry highly imprecise. Therefore, aperture photometry was executed using the Image Reduction and Analysis Facility (IRAF)⁹ software package *apphot*. The only exception to this analysis is PSR J1911-5958A, where unique *Hubble Space Telescope* (HST) observations required a different approach that will be discussed later.

We performed differential photometry using multiple comparison stars inversely weighted to their intrinsic variance over a single run of observation (a technique similar to that used in

Gilliland & Brown 1988 and Sokoloski et al. 2001). We ensured that all comparison stars are nonvariable by comparing each prospective star within the field of view with all others in every combination, and we used periodograms to ensure that no periodic trends are present.

We employed a variable-aperture algorithm to optimize the S/N of our photometry in order to deal with the varying PSFs. For each observation, a single representative frame was selected. For every star in the image we varied the size of the photometric aperture to find the size that optimized the S/N of the measurement. The optimum apertures were then scaled by the average FWHM of 2D Gaussians fitted to each star. For all other images the average FWHM of the stellar PSFs was measured and then scaled using the appropriate optimum scaling. Therefore, as the seeing fluctuated during an observation, the apertures were adjusted to be larger or smaller (a technique similar to that used in Deeg & Doyle 2001).

Once we have constructed the light curve for our target star, there are often long-timescale trends still present in the data, usually caused by the color/air-mass effect. We corrected for this effect by correlating our light curve with air mass and dividing a fitted trend from the light curve. For some observations this trend removal is insufficient, as other systematics may be present (e.g., variable sky background due to the moon). In these cases low-order polynomials (third-order or less) were used to remove this power.

3.2. Periodograms and Detection Limits

To detect pulsations in our light curve, we used a Lomb-Scargle periodogram (Scargle 1982), because our observations were frequently unevenly spaced and sometimes filled with large gaps. In the ZZ Ceti stars, low-amplitude pulsations show pulse shapes that are closely approximated by sinusoids, while high-amplitude pulsations frequently show narrow peaks in their shape (Fontaine & Brassard 2008). We feel that this justifies our use of a periodogram that is optimized for sinusoidal signals, because the pulsations that are “in the noise” are closely represented by sinusoids. High-amplitude pulsations, though

⁹IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under contract with the National Science Foundation; see <http://iraf.noao.edu>.

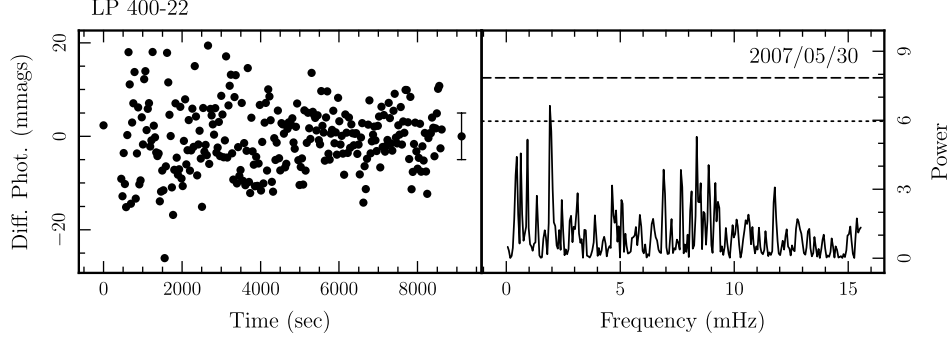


FIG. 11.—Light curve and Lomb-Scargle periodogram for LP 400-22 for observations taken at WIYN/OPTIC on 2007 May 30. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

nonsinusoidal, should still be easily detected, as they will likely be well above the noise.

We did not find any significant pulsations from He WDs. However, detection limits are still useful for both deciding if additional observations are required and for constraining theory. We measured our detection limits by using signal injection and the properties of a Lomb-Scargle periodogram. For a periodogram normalized by the variance of the total light curve, a false-alarm probability is naturally generated, which states that for M independent frequency bins, the probability that independent and normal noise produces power greater than z in any of the frequency bins is $\text{Prob}(\text{any power} > z) = 1 - (1 - e^{-z})^M$ (Scargle 1982; Horne & Baliunas 1986). Thus, a signal may be considered 90% significant if it has a false-alarm probability of 10%. However, our noise may not be precisely normal and

independent, due to potential unresolved pulsations and other systematic effects (such as atmospheric correlated noise) that may cause our data to be slightly correlated. This may cause us to slightly overestimate our detection limits. Our simulations of normal random data distributed with the same errors as our light curves have shown, however, that this effect is minimal. We thus determined our detection limits by injecting sinusoidal signals, integrated over the observation's exposure times, at a single frequency, and we tested all phases. We then increased the amplitude of the signal until 90% of all phases produce a peak in the periodogram at the injected frequency that is 90% significant. Finally, all frequencies accessible by the data were tested with this method. A low-frequency limit to the accessible frequencies was derived by testing the low-frequency content of the comparison stars where we do *not* correct for

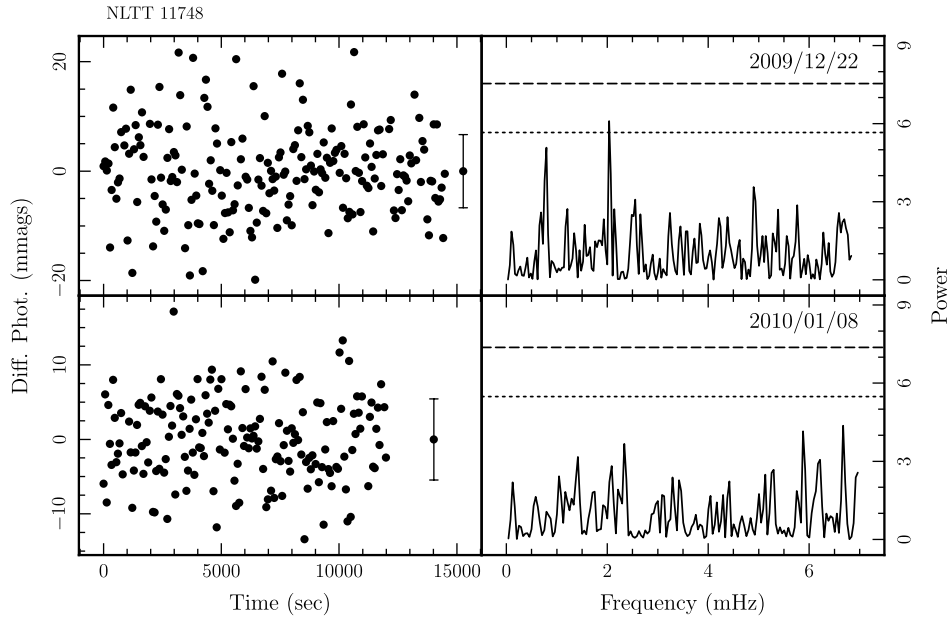


FIG. 12.—Light curves and Lomb-Scargle periodograms for NLTT 11748 for observations taken at FTN/Merope on 2009 December 22 and 2010 January 8. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

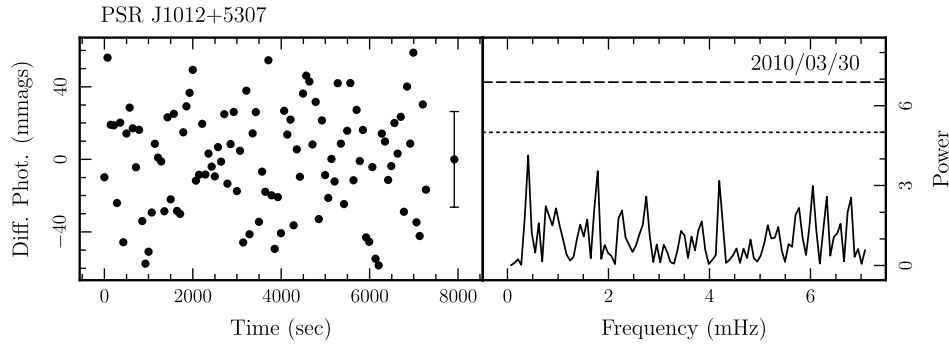


FIG. 13.—Light curve and Lomb-Scargle periodogram for PSR J1012 + 5307 for observations taken at WHT/ACAM on 2010 March 30. The dashed and dotted lines in the periodogram represent the 10% and 50% false-alarm probability thresholds given the number of frequency bins.

color/air-mass or other long-period trends. We report the low-frequency limit as the highest frequency where the red-noise power drops below the 50% false-alarm probability level. All frequencies below this limit will be affected by the color/air-mass and long-period trend corrections and are thus not reliable. For all observations, this low-frequency limit is a few times larger than the sampling limit of the inverse of the observation length. In all cases in this article, the detection limits were largely independent of frequency, allowing us to state one detection limit over a range of frequencies. The results of our observations are shown in Table 3 and are summarized in the next section and in Table 1.

4. SUMMARY OF OBSERVATIONS AND RESULTS

In the course of this article, several observing setups were employed. Here, we summarize the details of the telescopes and instrumentation used.

The William Herschel Telescope (WHT) is a 4.2 m telescope at the Roque de Los Muchachos Observatory on La Palma in the Canary Islands. We used the ACAM (auxiliary-port camera) imager and the SDSS g' filter with a circular $8.3'$ diameter field of view in a 2×2 pixel binning mode for an effective spatial resolution of $0.5'' \text{ pixel}^{-1}$. Using the fast-readout mode we maintained a total dead time of ≈ 11 s between exposures. All observations using this setup were carried out using service time, generally at bright time (near the calendar full moon).

The Wisconsin-Indiana-Yale-NOAO (WIYN) telescope is a 3.5 m telescope at the Kitt Peak National Observatory in Arizona. We used the Orthogonal Parallel Transfer Imaging Camera (OPTIC) imager (Howell et al. 2003) and the broadband filter BG-39 ($\lambda_c \approx 480$ nm, FWHM ≈ 260 nm) with a $9.5' \times 9.5'$ field of view in a 2×2 pixel binning mode, for an effective spatial resolution of $0.28'' \text{ pixel}^{-1}$. The total dead time between exposures was ≈ 10 s. All observations using this setup were carried out in a single observing run that occurred in late 2007 May and early 2007 June.

The Palomar 60 inch (P60) telescope is a 1.5 m robotically controlled telescope at the Palomar Observatory on Mount

Palomar in southern California. We used the P60CCD imager and the Gunn g filter with a $12.9' \times 12.9'$ field of view with a spatial resolution of $0.38'' \text{ pixel}^{-1}$. We decreased readout times by reading out half of the two CCD chips for a total dead time between exposures of ≈ 20 and 40 s. The large variation in dead time was due to a software issue in the telescope software. This setup was used in late 2006 and early 2007; however, most of the data taken during that time were for other projects or were too degraded by weather.

The Faulkes Telescope North (FTN) is a 2.0 m robotically controlled telescope on Haleakala, Hawaii, and is part of the Las Cumbres Observatory Global Telescope.¹⁰ We used the Merope imager and the SDSS g' filter with a $4.75' \times 4.75'$ field of view in a 2×2 pixel binning mode for an effective spatial resolution of $0.28'' \text{ pixel}^{-1}$. The total dead time between exposures was ≈ 22 s. This setup was used in late 2009 December and early 2010 January.

We used the Wide Field Camera 3 (WFC3) ultraviolet imaging spectrograph (UVIS) channel and the F200LP filter ($\lambda_c \approx 490$ nm, FWHM ≈ 250 nm) on *HST*. To decrease readout times and to avoid onboard memory-buffer overflow, we used the UVIS1-C512A-SUB subarray, a 512×512 pixel subsection of the CCD. This gave us a $20.5'' \times 20.5''$ field of view and a total dead time of 55 s. This setup was used in 2010 March for four orbits (6 hr).

Most of our observations and objects were routine, and the results are well summarized in Tables 1, 2, and 3 and Figures 2–14. However, a few of our objects and observations require special discussion.

SDSS J1330 + 0127.—was observed once using the WIYN/OPTIC setup. No pulsations were detected to a limit of 44 mmag over a frequency range of 0.3–5 mHz. This observation was plagued by intermittent cirrus clouds and a near-full moon creating a high sky background. This object has a measured $\log g = 5.56 \pm 0.16$ and $T_{\text{eff}} = 8,091 \pm 75$ K (Eisenstein et al. 2006). The SDSS spectra were likely contaminated by a companion

¹⁰ See <http://www.lcogt.net>.

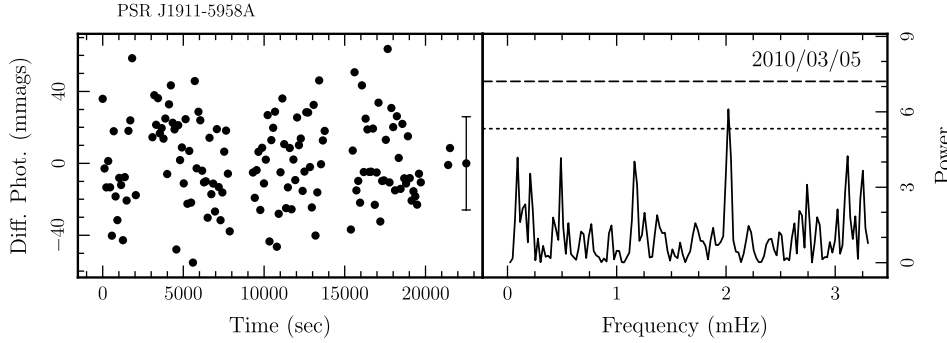


FIG. 14.—Light curve and Lomb-Scargle periodogram for PSR J1911-5958A for observations taken at *HST*/WFC3 on 2010 March 3. The dashed and dotted lines in the periodogram represent the 10% (and 50%) false-alarm probability thresholds given the number of frequency bins.

star's flux. A reanalysis by Silvestri et al. (2006) measured $\log g = 8.1 \pm 0.2$ and $T_{\text{eff}} = 10,617 \pm 452$ K. This object is likely not a He WD.

SDSS J1435 + 3733.—is a partially eclipsing binary system of a C/O WD and M-dwarf discovered by Steinfadt et al. (2008a) using the FTN/Merope setup. All frames in eclipse were removed for pulsation analysis. No pulsations were detected to a limit of 4 mmag over a frequency range of 0.9–17 mHz. This object has a measured $\log g = 6.86 \pm 0.04$ and $T_{\text{eff}} = 11,062 \pm 58$ K (Eisenstein et al. 2006). The SDSS spectra were definitely contaminated by a companion star's flux (Steinfadt et al. 2008a; Rebassa-Mansergas et al. 2007; Pyrzas et al. 2009). A reanalysis by Rebassa-Mansergas et al. (2007) and Pyrzas et al. (2009) measured $\log g = 7.62 \pm 0.12$ and $T_{\text{eff}} = 12,536 \pm 488$ K, which gives a derived model mass of $0.41 M_{\odot}$, near the C/O WD and He WD composition boundary.

SDSS J2049 + 0005.—was observed once using the WIYN/OPTIC setup. No pulsations were detected to a limit of 21 mmag over a frequency range of 0.4–7 mHz. This observation was plagued by intermittent cirrus clouds and a near-full moon creating a high sky background. This object has a measured $\log g = 5.48 \pm 0.10$ and $T_{\text{eff}} = 8660 \pm 144$ K (Eisenstein et al. 2006). Eisenstein et al. (2006) noted multiple minima in their model fits, as well as discrepant SDSS g' -band flux compared with best-fit models, and suggested that a $\log g < 5.0$ was likely. Kilic et al. (2007a) confirmed this finding and concluded that this star is likely a distant A2 star.

SDSS J2240-0935.—was observed 4 times using the P60/P60CCD setup. No pulsations were detected to a limit of 14 mmag over 0.6–7 mHz. This object has a measured $\log g = 6.96 \pm 0.09$ and $T_{\text{eff}} = 11,449 \pm 205$ K (Eisenstein et al. 2006). The SDSS spectra were likely contaminated by a companion star's flux. A reanalysis by Heller et al. (2009) measured $\log g = 7.5$ and $T_{\text{eff}} = 13,000$ K.

NLTT 11748.—is the first totally eclipsing detached double white dwarf binary system discovered by Steinfadt et al. (2010b) using the FTN/Merope setup. All frames in eclipse were removed for pulsation analysis. No pulsations were

detected to a limit of 4 mmag over 0.2–7 mHz. This object has a measured $\log g = 6.54 \pm 0.05$ and $T_{\text{eff}} = 8690 \pm 140$ K (Kilic et al. 2010a) with a derived mass of $0.18 M_{\odot}$. Shporer et al. (2010) found direct evidence for Doppler beaming in a long-timescale photometric survey. The Doppler beaming signal was removed in the low-order polynomial division in the reduction and should only affect frequencies less than 0.1 mHz (half the orbital period of 5.6 hr). In the future, eclipse and radial velocity analysis will produce a high-precision measurement of the masses and radii of both components in this object. Since its discovery, three additional eclipsing He WD binary systems have been discovered (Pyrzas et al. 2011; Brown et al. 2011b; Parsons et al. 2011).

PSR J1012 + 5307.—was observed one time using the WHT/ACAM setup. No pulsations were detected to a limit of 20 mmag over the frequency range of 0.2–7 mHz. This object has a measured $\log g = 6.34 \pm 0.20$ and $T_{\text{eff}} = 8670 \pm 300$ K with a derived model mass of $0.18 M_{\odot}$ (Callanan et al. 1998) and also $\log g = 6.75 \pm 0.07$ and $T_{\text{eff}} = 8550 \pm 25$ K with a derived model mass of $0.21 M_{\odot}$ (van Kerkwijk et al. 1996). Neither of these measurements is obviously superior.

PSR J1911-5958A.—was observed one time using the *HST*/WFC3 setup over four orbits. The photometric analysis for this object was uniquely different from all other objects. Due to crowding in the field (PSR J1911-5958A is in a globular cluster), as well as the unique PSF of *HST*, aperture photometry is not the most precise method. PSF-fitting photometry is ideal, since the PSF of the WFC3 instrument is well understood (J. Anderson 2010, private communication). We used the PSF-fitting software developed by J. Anderson (similar to Anderson & King 2006) to extract the photometry from our data. Since most atmospheric and sky effects are nonexistent in space, variable-aperture photometry was not required or used. We used 36 comparison stars with magnitudes between 17.5 and 25.0 to ensure the quality of our photometry. Due to a stray-light artifact, our background was increased by $\approx 30\%$. No pulsations were detected to a limit of 16 mmag over the frequency range of 0.05–3.25 mHz. This object has a measured

$\log g = 6.44 \pm 0.20$ and $T_{\text{eff}} = 10,090 \pm 150$ K with a derived model mass of $0.18 M_{\odot}$ (Bassa et al. 2006a).

5. CONCLUSIONS

Of the 12 objects we observed in this article, seven are likely very low mass ($<0.2 M_{\odot}$) He WDs. Whether or not these observations offer strong constraints on theory is currently an open matter. First we must consider the accuracy of the $\log g$ and T_{eff} measurements that place these objects within the predicted pulsation parameter space. Unfortunately, very low mass WD atmospheres are not well understood. Due to their low surface temperatures ($T_{\text{eff}} \lesssim 9000$ K), the H atmospheres are not sufficiently ionized, and neutral broadening of the H absorption lines plays an important role. The theory of neutral broadening is not well understood and may account for significant errors in line widths (Barklem et al. 2000; Allard et al. 2004; Tremblay et al. 2010; D. Koester 2010, private communication). As such, the reported errors on $\log g$ and T_{eff} for these objects are always the statistical error of the fit to theoretical atmosphere grids and do not take into account the uncertainties of the theory itself, which can increase the uncertainty by as much as a factor of 2 or more (D. Koester 2010, private communication). Further, the uncertainty of the atmosphere theory is likely not random, but rather a shift in a specific direction in the $T_{\text{eff}}\text{--}\log g$ plane, affecting the parameter space location of all very low mass He WDs. This is why NLTT 11,748 is such an important eclipsing binary system (Steinfadt et al. 2010b). As a potential double-lined spectroscopic eclipsing binary, it could offer high-precision *model-independent* determinations of both the mass and radius of the He WD component, providing a high-precision constraint on its gravity for which any atmospheric gravity measurement must match.

Of added concern is the location of the instability strip blue edge itself. Steinfadt et al. (2010a) derived the blue edge to be where pulsation modes of eigenvalues $\ell = 1$ and $n = 1$ were driven. This offers a temperature upper limit to the $\ell = 1$ instability strip, because the blue edge shifts redward when higher radial orders of n are considered. In the few cases where pulsation-mode eigenvalues have been identified observationally (e.g., HL Tau 76 [Pech et al. 2006] and G117-B15A and R548 [Bradley 1998]) the highest-amplitude observed modes range in n from 1 to several tens. If very low mass He WDs preferentially only excite the

higher radial order modes, then the blue edge could shift redward by a few hundred kelvins, potentially removing our three best candidates from the potential instability region. Uncertainties in the precise convective efficiency and the lack of a detailed non-adiabatic analysis could further shift this blue edge by another several hundred kelvins in either direction (Winget et al. 1982a; Fontaine & Brassard 2008).

An additional observational concern is whether our detection limits are truly meaningful. Only NLTT 11748, LP 400-20, SDSS J0822 + 2743, and SDSS J1435 + 3733 met our fiducial pulsation limit of 10 mmag, and only NLTT 11748, SDSS J1435 + 3733, and SDSS J2240-0935 were observed more than once. It is well known that pulsation amplitudes can be small, 3–4 mmag, and vary in detectability from one night to the next (Mukadam et al. 2006; Steinfadt et al. 2008b). Therefore, additional observations are required for many of our objects, to both tighten their detection limits and gain multiple epochs of observation so that these null detections can be even more convincing.

Finally, the theoretical analysis by Steinfadt et al. (2010a) involved only an extension of the Brickhill convective driving criterion using adiabatic assumptions that may not be valid at the lower gravities of the very low mass He WDs. Further, very little theoretical work has been done between 0.2 and $0.45 M_{\odot}$ (Althaus et al. 2004; Arras et al. 2006). To address these concerns, an extension of the nonadiabatic driving work by Brickhill (1991) and Wu & Goldreich (1999) should be used to firmly predict the location of the He WD instability strip.

We thank both referees for comments that improved our article. This work was supported by the National Science Foundation under grants PHY 05-51164, AST 07-07633, and AST 11-09174. D. L. K. was partially supported by NASA through Hubble Fellowship Grant 01207.01-A awarded by the Space Telescope Science Institute, which is operated by AURA, Inc., for NASA, under contract NAS 5-26555. This article uses observations obtained with facilities of the Las Cumbres Observatory Global Telescope. Support for Program number HST-GO-11581.01-A was provided by NASA through a grant from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Incorporated, under NASA contract NAS5-26555.

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